

# Stability of the Regular Hayward Thin-Shell Wormholes

M. Sharif <sup>\*</sup> and Saadia Mumtaz <sup>†</sup>

Department of Mathematics, University of the Punjab,  
Quaid-e-Azam Campus, Lahore-54590, Pakistan.

## Abstract

The aim of this paper is to construct regular Hayward thin-shell wormholes and analyze their stability. We adopt Israel formalism to calculate surface stresses of the shell and check the null and weak energy conditions for the constructed wormholes. It is found that the stress-energy tensor components violate the null and weak energy conditions leading to the presence of exotic matter at the throat. We analyze the attractive and repulsive characteristics of wormholes corresponding to  $a^r > 0$  and  $a^r < 0$ , respectively. We also explore stability conditions for the existence of traversable thin-shell wormholes with arbitrarily small amount of fluids describing cosmic expansion. We find that the spacetime has non-physical regions which give rise to event horizon for  $0 < a_0 < 2.8$  and the wormhole becomes non-traversable producing a black hole. The non-physical region in the wormhole configuration decreases gradually and vanishes for the Hayward parameter  $l = 0.9$ . It is concluded that the Hayward and Van der Waals quintessence parameters increase the stability of thin-shell wormholes.

**Keywords:** Thin-shell wormholes; Israel formalism; Stability.

**PACS:** 04.20.-q; 04.40.Nr; 04.40.Gz; 04.70.Bw.

---

<sup>\*</sup>msharif.math@pu.edu.pk

<sup>†</sup>sadiamumtaz17@gmail.com

# 1 Introduction

One of the most interesting attributes of general relativity is the possible existence of hypothetical geometries having non-trivial topological structure. Misner and Wheeler [1] described these topological features of spacetime as solutions of the Einstein field equations known as wormholes. A “wormhole” having a tunnel with two ends allows a short way associating distant regions of the universe. Besides the lack of some observational evidences, wormholes are regarded as a part of black holes (BH) family [2]. The simplest example is the Schwarzschild wormhole that connects one part of the universe to other through a bridge. This wormhole is not traversable as it does not allow a two way communication between two regions of the spacetime leading to the contraction of wormhole throat.

Physicists have been motivated by the proposal of Lorentzian traversable wormholes given by Morris and Thorne [3]. In case of traversable wormholes, the wormhole throat is threaded by exotic matter which causes repulsion against the collapse of wormhole throat. The most distinguishing property of these wormholes is the absence of event horizon which enables observers to traverse freely across the universe. It was shown that a BH solution with horizons could be converted into wormhole solution by adding some exotic matter which makes the wormhole stable [4]. Traversable wormhole solutions must satisfy the flare-out condition preserving their geometry due to which the wormhole throat remains open. The existence of exotic matter yields the violation of null (NEC) and weak energy conditions (WEC) which is the basic property for traversable wormholes. Null energy condition is the weakest one whose violation gives rise to the violation of WEC and strong energy conditions (SEC). The exotic matter is characterized by stress-energy tensor components determined through Israel thin-shell formalism [5].

Thin-shell wormholes belong to one of the wormhole classes in which exotic matter is restricted at the hypersurface. To make sure that the observer does not encounter non-physical zone of BH, a thin-shell strengthens the wormhole provided that it has an exotic matter for its maintenance against gravitational collapse. The physical viability of thin-shell wormholes is a debatable issue due to inevitable amount of exotic matter which is an essential ingredient for the existence as well as stability of wormholes. The amount of exotic matter can be quantified by the volume integral theorem which is consistent with the concept that a small quantity of exotic matter is needed to support wormhole [6]. Visser [7] developed an elegant technique of cut

and paste to minimize the amount of exotic matter by restricting it at the edges of throat in order to obtain a more viable thin-shell wormhole solution.

It is well-known that thin-shell wormholes are of significant importance if they are stable. The stable/unstable wormhole models can be investigated either by applying perturbations or by assuming equation of state (EoS) supporting exotic matter at the wormhole throat. In this context, many authors constructed thin-shell wormholes following Visser's cut and paste procedure and discussed their stability. Kim and Lee [8] investigated stability of charged thin-shell wormholes and found that charge affects stability without affecting the spacetime itself. Ishak and Lake [9] analyzed stability of spherically symmetric thin-shell wormholes. Lobo and Crawford [10] studied spherically symmetric thin-shell wormholes with cosmological constant ( $\Lambda$ ) and found that stable solutions exist for positive values of  $\Lambda$ . Eiroa and Romero [11] studied linearized stability of charged spherical thin-shell wormholes and found that presence of charge significantly increases the possibility of stable wormhole solutions. Sharif and Azam [12] explored both stable and unstable configurations for spherical thin-shell wormholes. Sharif and Mumtaz analyzed stable wormhole solutions from regular ABG [13] and ABGB [14] spacetimes in the context of different cosmological models for exotic matter.

It is found that one may construct a traversable wormhole theoretically with arbitrarily small amount of fluids describing cosmic expansion. In order to find any realistic source for exotic matter, different candidates of dark energy have been proposed like tachyon matter [15], family of Chaplygin gas [16, 17], phantom energy [18] and quintessence [19]. Eiroa [20] assumed generalized Chaplygin gas to study the dynamics of spherical thin-shell wormholes. Kuhfittig [21] analyzed stability of spherical thin-shell wormholes in the presence of  $\Lambda$  and charge by assuming phantom like EoS at the wormhole throat. Sharif and collaborators discussed stability analysis of Reissner-Nordström [22] and Schwarzschild de Sitter as well as anti-de Sitter [23] thin-shell wormholes in the vicinity of generalized cosmic Chaplygin gas (GCCG) and modified cosmic Chaplygin gas (MCCG). Some physical properties of spherical traversable wormholes [24] as well as stability of cylindrical thin-shell wormholes [25, 26] have been studied in the context of GCCG, MCCG and Van der Waals (VDW) quintessence EoS. Recently, Halilsoy *et al.* [27] discussed stability of thin-shell wormholes from regular Hayward BH by taking linear, logarithmic and Chaplygin gas models and found stable solutions for increasing values of Hayward parameter.

This paper is devoted to construct thin-shell wormholes from regular Hayward BH by considering three different models of exotic matter at the throat. The paper is organized as follows. In section 2, we construct regular Hayward thin-shell wormholes and analyze various physical aspects of these constructed thin-shell wormholes. Section 3 deals with stability formalism of the regular Hayward thin-shell wormholes in the vicinity of VDW quintessence EoS and Chaplygin gas models. We find different throat radii numerically and show their expansion or collapse with different values of parameters. Finally, we provide summary of the obtained results in the last section.

## 2 Regular Hayward Black Hole and Wormhole Construction

The static spherically symmetric regular Hayward BH [28] is given by

$$ds^2 = -F(r)dt^2 + F^{-1}(r)dr^2 + G(r)(d\theta^2 + \sin^2\theta d\phi^2), \quad (1)$$

where  $G(r) = r^2$  and  $F(r) = 1 - \frac{2Mr^2}{r^3 + 2Ml^2}$ ,  $M$  and  $l$  are positive constants. This regular BH is chosen for thin-shell wormhole because a regular system can be constructed from a finite energy and its evolution is more acceptable. This reduces to de Sitter BH for  $r \rightarrow 0$ , while the metric function for the Schwarzschild BH is obtained as  $r \rightarrow \infty$ . Its event horizon is the largest root of the equation

$$r^3 - 2Mr^2 + 2Ml^2 = 0. \quad (2)$$

This analysis of the roots shows a critical ratio  $\frac{l}{M_*} = \frac{4}{3\sqrt{3}}$  and radius  $r_* = \sqrt{3}l$  such that for  $r > 0$  and  $M < M_*$ , the given spacetime has no event horizon yielding a regular particle solution. The regular Hayward BH admits a single horizon if  $r = r_*$  and  $M = M_*$ , which represents a regular extremal BH. At  $r = r_{\pm}$  and  $M > M_*$ , the given spacetime becomes a regular non-extremal BH with two event horizons.

We implement the standard cut and paste procedure to construct a time-like thin-shell wormhole. In this context, the interior region of the regular Hayward BH is cut with  $r < a$ . The two 4D copies are obtained which are glued at the hypersurface  $\Sigma^{\pm} = \Sigma = \{r = a\}$ . Infact this technique treats the hypersurface  $\Sigma$  as the minimal surface area called wormhole throat. The exotic matter is concentrated at the hypersurface making the wormhole solution a thin-shell. We can take coordinates  $\chi^i = (\tau, \theta, \phi)$  at the shell. The

induced 3D metric at  $\Sigma$  with throat radius  $a = a(\tau)$  is defined as

$$ds^2 = -d\tau^2 + a^2(\tau)(d\theta^2 + \sin^2\theta d\phi^2), \quad (3)$$

where  $\tau$  is the proper time on the shell.

This construction requires the fulfillment of flare-out condition by the throat radius  $a$ , i.e., the embedding function  $G(r)$  in Eq.(1) should satisfy the relation  $G'(a) = 2a > 0$ . The thin layer of matter on  $\Sigma$  causes the extrinsic curvature discontinuity. In this way, Israel formalism is applied for the dynamical evolution of thin-shell which allows matching of two regions of spacetime partitioned by  $\Sigma$ . We find non-trivial components of the extrinsic curvature as

$$K_{\tau\tau}^{\pm} = \mp \frac{F'(a) + 2\ddot{a}}{2\sqrt{F(a) + \dot{a}^2}}, \quad K_{\theta\theta}^{\pm} = \pm a\sqrt{F(a) + \dot{a}^2}, \quad K_{\phi\phi}^{\pm} = \alpha^2 K_{\theta\theta}^{\pm}, \quad (4)$$

where dot and prime stand for  $\frac{d}{d\tau}$  and  $\frac{d}{dr}$ , respectively. To determine surface stresses at the shell, we use Lanczos equations, which are the Einstein equations given by

$$S_{ij} = \frac{1}{8\pi} \{g_{ij}K - [K_{ij}]\}, \quad (5)$$

where  $[K_{ij}] = K_{ij}^+ - K_{ij}^-$  and  $K = tr[K_{ij}] = [K^i_i]$ . The surface energy-momentum tensor  $S_{ij}$  yields the surface energy density  $S_{\tau\tau} = \sigma$  and surface pressures  $S_{\theta\theta} = p = S_{\phi\phi}$ . Solving Eqs.(4) and (5), the surface stresses are obtained as

$$\sigma = -\frac{1}{2\pi a} \sqrt{F(a) + \dot{a}^2}, \quad (6)$$

$$p = p_{\theta} = p_{\phi} = \frac{1}{8\pi} \frac{2\dot{a}^2 + 2a\ddot{a} + 2F(a) + aF'(a)}{a\sqrt{F(a) + \dot{a}^2}}. \quad (7)$$

In order to prevent contraction of wormhole throat, matter distribution of surface energy-momentum tensor must be negative which indicates the existence of exotic matter making the wormhole traversable [7]. The amount of this matter should be minimized for the sake of viable wormhole solutions. We note from Eqs.(6) and (7) that  $\sigma < 0$  and  $\sigma + p < 0$  showing the violation of NEC and WEC for different values of  $M$ ,  $l$  and  $a$ . In Figure 1, we plot a graph for pressure showing that pressure is a decreasing function of the throat radius (left hand side), while the other graph shows violation of energy conditions associated with regular Hayward thin-shell wormholes.

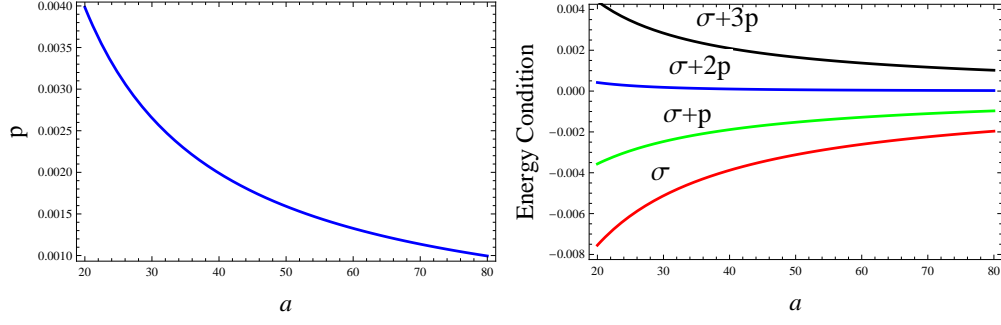


Figure 1: Plots of  $p$  and energy conditions with  $M = 1$  and  $l = 0.9$ .

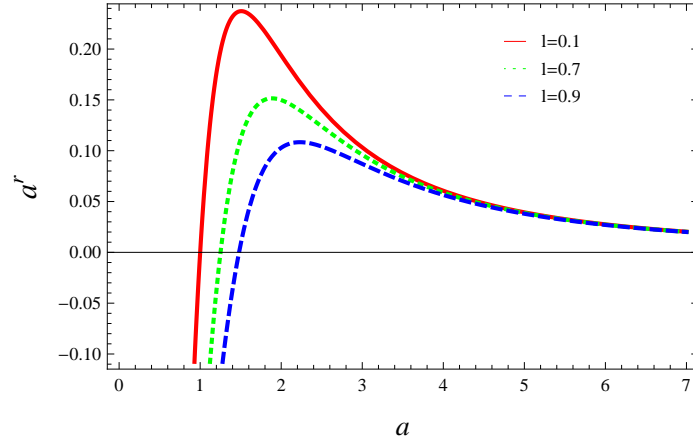


Figure 2: Plots of  $a^r$  with  $M = 1$  and  $l = 0.1, 0.7, 0.9$ . The wormhole is attractive for  $a^r > 0$  and repulsive for  $a^r < 0$ .

Now we explore the attractive and repulsive characteristics of the regular Hayward thin-shell wormholes. In this context, we need to compute the observer's four-acceleration

$$a^\mu = u^\mu_{;\nu} u^\nu,$$

where  $u^\mu = \frac{dx^\mu}{d\tau} = (\frac{1}{\sqrt{F(r)}}, 0, 0, 0)$  is the observer's four-velocity. The non-zero four-acceleration component corresponding to the given spacetime is calculated as

$$a^r = \Gamma^r_{tt} \left( \frac{dt}{d\tau} \right)^2 = \frac{Mr^4 - 4M^2 l^2 r}{(r^3 + 2ml^2)^2}, \quad (8)$$

for which the geodesic equation has the following form

$$\frac{d^2 r}{d\tau^2} = -\Gamma^r_{tt} \left( \frac{dt}{d\tau} \right)^2 = -a^r.$$

An important condition for traversing through wormhole implies that an observer should not be dragged away by enormous tidal forces. It is required that the acceleration felt by observer must not exceed the Earth's acceleration. It is worth stressing here that a wormhole will be attractive in nature if its radial acceleration is positive, i.e.,  $a^r > 0$ . This supports the fact that an observer must have an outward-directed radial acceleration  $a^r$  in order to keep away from being pulled by the wormhole. On the other hand, it will exhibit repulsive characteristics for  $a^r < 0$ . In this case, an observer must move with an inward directed radial acceleration to avoid being repelled by the wormhole. The attractive and repulsive characteristics of the regular Hayward thin-shell wormholes are shown in Figure 2.

Some researchers are excited by the possibility of wormholes in reality. It appears feasible to keep the wormhole throat open long enough such that an object can traverse easily through it if the throat is threaded by exotic matter. The total amount of exotic matter is quantified by the integral theorem [6]

$$\Omega = \int [\rho + p_r] \sqrt{-g} d^3x. \quad (9)$$

By introducing radial coordinate  $R = r - a$ , we have

$$\Omega = \int_0^{2\pi} \int_0^\pi \int_{-\infty}^{+\infty} [\rho + p_r] \sqrt{-g} dR \sin \theta d\theta d\phi. \quad (10)$$

The wormhole shell, being thin, does not apply any pressure leading to  $p_r = 0$ . Using  $\rho = \delta(R)\sigma(a)$ , we have

$$\Omega_a = \int_0^{2\pi} [\rho\sqrt{-g}]|_{r=a} d\phi = 2\pi a\sigma(a). \quad (11)$$

Inserting the value of surface energy density  $\sigma(a)$ , the above expression yields

$$\Omega_a = -\sqrt{\frac{a_0^3 + 2Ml^2 - 2Ma_0^2}{a_0^3 + 2Ml^2}}, \quad (12)$$

where  $a_0$  is the wormhole throat radius. It is interesting to note that construction of a traversable wormhole is possible theoretically with vanishing amount of exotic matter. This amount can be made infinitesimally small by choosing exotic fluids explaining cosmic expansion.

### 3 Stability of Thin-Shell Wormholes

Here we study the formation of thin-shell wormholes from regular Hayward BH and analyze their stability under linear perturbations. The surface energy density and surface pressure corresponding to static wormhole configuration at  $a = a_0$  become

$$\sigma_0 = -\frac{\sqrt{F(a_0)}}{2\pi a_0}, \quad p_0 = \frac{1}{8\pi} \frac{2F(a_0) + a_0 F'(a_0)}{a_0 \sqrt{F(a_0)}}. \quad (13)$$

The energy density and pressure follow the conservation identity  $S^{ij}_{;j} = 0$ , which becomes for the line element (1) as

$$\frac{d}{d\tau}(\sigma\Phi) + p\frac{d\Phi}{d\tau} = 0, \quad (14)$$

where  $\Phi = 4\pi a^2$  corresponds to wormhole throat area. Using  $\sigma' = \frac{\dot{\sigma}}{a}$ , the above equation can be written as

$$a\sigma' = -2(\sigma + p). \quad (15)$$

The thin-shell equation of motion can be obtained by rearranging Eq.(6) as  $\dot{a}^2 + \Psi(a) = 0$ , which determines wormhole dynamics while the potential function  $\Psi(a)$  is defined by

$$\Psi(a) = F(a) - [2\pi a\sigma(a)]^2. \quad (16)$$



In order to explore wormhole stability, we expand  $\Psi(a)$  around  $a = a_0$  using Taylor's series as

$$\Psi(a) = \Psi(a_0) + \Psi'(a_0)(a - a_0) + \frac{1}{2}\Psi''(a_0)(a - a_0)^2 + O[(a - a_0)^3]. \quad (17)$$

The first derivative of Eq.(16) through (15) takes the form

$$\Psi'(a) = F'(a) + 8\pi^2 a \sigma(a) [\sigma(a) + 2p(a)]. \quad (18)$$

The stability of wormhole static solution depends upon  $\Psi''(a_0) \geq 0$  and  $\Psi'(a_0) = 0 = \Psi(a_0)$ . The surface stresses for static configuration (13) yield

$$\sigma_0 = -\frac{\sqrt{a_0^3 + 2Ml^2 - 2Ma_0^2}}{2\pi a_0 \sqrt{a_0^3 + 2Ml^2}}, \quad p_0 = \frac{a_0^3 + 2Ml^2 - 4Ma_0^2}{4\pi a_0 \sqrt{(a_0^3 + 2Ml^2)(a_0^3 + 2Ml^2 - 2Ma_0^2)}}. \quad (19)$$

The choice of model for exotic matter has significant importance in the dynamical investigation of thin-shell wormholes. In a recent work, Halilsoy *et al.* [27] examined the dynamics of Hayward thin-shell wormholes for linear, logarithmic and Chaplygin gas models. In this paper, we take VDW quintessence, GCCG and MCGG fluids at the shell to study stability of regular Hayward thin-shell wormholes. We will explore the possibility for existence of stable traversable wormhole solutions by taking different EoS for exotic matter. In the following, we adopt standard stability formalism by in the context of the above candidates of dark energy as exotic matter.

### 3.1 Van der Waals Quintessence

Firstly, we model the exotic matter by VDW quintessence EoS which is a remarkable scenario to describe accelerated expansion of the universe without the presence of exotic fluids. The EoS for VDW quintessence is given by

$$p = \frac{\gamma\sigma}{1 - B\sigma} - \alpha\sigma^2, \quad (20)$$

where  $\alpha$ ,  $B$  and  $\gamma$  are EoS parameters. The specific values of these parameters lead to accelerated and decelerated periods. Inserting Eq.(13) in (20), the equation for static configuration is obtained as

$$\left\{ a_0^2 \psi'(a_0) + 2a_0 \psi(a_0) + \frac{2\alpha}{\pi} [\psi(a_0)]^{\frac{3}{2}} \right\} \left\{ 2\pi^2 a_0 + B\pi \sqrt{\psi(a_0)} \right\} + 2\gamma(2\pi a_0)^2 \psi(a_0) = 0. \quad (21)$$

The EoS turns out to be

$$\sigma'(a) + 2p'(a) = \sigma'(a) \left\{ 1 + \frac{2}{1 - B\sigma(a)} [\gamma - 2\alpha\sigma(a) + B\{p(a) + 3\alpha\sigma^2(a)\}] \right\}. \quad (22)$$

It is found that  $\Psi(a) = \Psi'(a) = 0$  by substituting the values of  $\sigma(a_0)$  and  $p(a_0)$ , while the second derivative of  $\Psi$  through Eqs.(18) and (22) becomes

$$\begin{aligned} \Psi''(a_0) = & F''(a_0) + \frac{[F'(a_0)]^2}{2F(a_0)} \left[ \frac{B\sqrt{F(a_0)}}{2\pi a_0 + \sqrt{F(a_0)}} - 1 \right] + \frac{F'(a_0)}{a_0} \left[ 1 \right. \\ & + \frac{1}{2\pi a_0 + \sqrt{F(a_0)}} \left\{ 4\pi a_0 \gamma + 4\alpha \sqrt{F(a_0)} + B \left( 2\sqrt{F(a_0)} \right. \right. \\ & + \left. \left. \frac{3\alpha B \sqrt{F(a_0)}}{\pi a_0} \right) \right\} \left. \right] - \frac{2F(a_0)}{a_0^2} (1 + \gamma) \left[ 1 + \frac{1}{2\pi a_0 + \sqrt{F(a_0)}} \right. \\ & \times \left. \left\{ 4\pi a_0 \gamma + 4\alpha \sqrt{F(a_0)} + B \left( 2\sqrt{F(a_0)} + \frac{3\alpha B \sqrt{F(a_0)}}{\pi a_0} \right) \right\} \right]. \end{aligned} \quad (23)$$

Now we formulate static solutions for which the dynamical equation through Eq.(21) takes the form

$$\begin{aligned} & 2a_0^4(a_0^3 + Ml^2 - Ma_0^2) + 8M^2l^2a_0^2(1 - 2a_0^2) + \frac{2\alpha}{\pi} \left[ 2\pi^2a_0(a_0^3 + 2Ml^2)^{\frac{1}{2}} \right. \\ & \times (a_0^3 + Ml^2 - 2Ma_0^2)^{\frac{3}{2}} + B\pi(a_0^3 + 2Ml^2 - 2Ma_0^2)^2 \left. \right] + 2\gamma(2\pi a_0^2)^2 \\ & \times [a_0^3(a_0^3 + 4Ml^2 - 2Ma_0^2) + 4Ml^2a_0^2(a_0 - 1)] = 0, \end{aligned} \quad (24)$$

whose solutions correspond to static Hayward thin-shell wormholes. In order to explore the wormhole stability, we evaluate numerical value of throat radius  $a_0$  (static) from Eq.(24) and substitute in Eq.(23). We choose Hayward parameter  $l = 0, 0.1, 0.7, 0.9$  and check the role of increasing values of  $l$  on the stability of Hayward thin-shell wormholes. We are interested to find the possibility for existence of traversable thin-shell wormholes and to check whether the wormhole throat will expand or collapse under perturbation. For the existence of static stable solutions,  $\Psi'' > 0$  and  $a_0 > r_h$ , while  $\Psi'' < 0$  and  $a_0 > r_h$  hold for unstable solutions. For  $a_0 \leq r_h$ , no static solution exists leading to non-physical region (grey zone). In this region, the stress-energy

tensor may vanish leading to an event horizon which makes the wormhole no more traversable. The stable and unstable solutions correspond to green and yellow zones, respectively. The graphical results in Figures 3-4 can be summarized as follows.

For  $\gamma \in (-\infty, -0.3]$ , only unstable solutions exist corresponding to  $l = 0, 0.1$ , while both (stable and unstable) wormhole configurations appear by increasing the values of Hayward parameter, i.e.,  $l = 0.7, 0.9$  as shown in Figure 3. For  $l = 0.7$ , the wormhole is initially stable but its throat continues to expand leading to unstable solution. The non-physical region in the wormhole configuration decreases gradually and vanishes for  $l = 0.9$ . In this case, we find unstable wormhole solution for  $a_0 < 2$  which leads to the collapse of wormhole throat as  $B\alpha^{-(1+\gamma)}$  approaches its maximum value. For increasing  $\gamma$ , i.e.,  $\gamma \in [0.1, 0.9]$ , there exist unstable and stable configurations.

Finally, we examine the stability of Hayward thin-shell wormholes for  $\gamma \in [1, \infty)$  and find only stable solutions with  $l = 0, 0.1, 0.7$  which shows the expansion and traversability of wormhole throat. The unstable solution also appears for  $l = 0.9$  and  $a_0 < 2$  leading to non-traversable wormhole due to its collapse. We find stable solutions for  $a_0 > 2$  and the wormhole throat expands which let the wormhole to open its mouth. Figure 4 shows the corresponding results for  $\gamma = 1$ . This graphical analysis shows that the wormhole exhibits physical regions (stable/unstable) corresponding to different values of Hayward parameter. Since the regular Hayward wormholes are singularity free due to their regular centers but the spacetime has event horizons which give rise to non-physical regions for  $0 < a_0 < 2.8$  and makes the wormhole non-traversable. We find that the wormhole can be made traversable as well as stable by tuning the Hayward parameter to its large value. Also, it is noted that  $\gamma = 1$  is the most fitted value to analyze only stable solutions.

### 3.2 Generalized Cosmic Chaplygin Gas

Now we assume GCCG [29] to support the exotic matter at the shell. Chaplygin gas is a hypothetical substance that satisfies an exotic EoS. The EoS for GCCG is defined as

$$p = -\frac{1}{\sigma^\gamma} [E + (\sigma^{1+\gamma} - E)^{-w}], \quad (25)$$

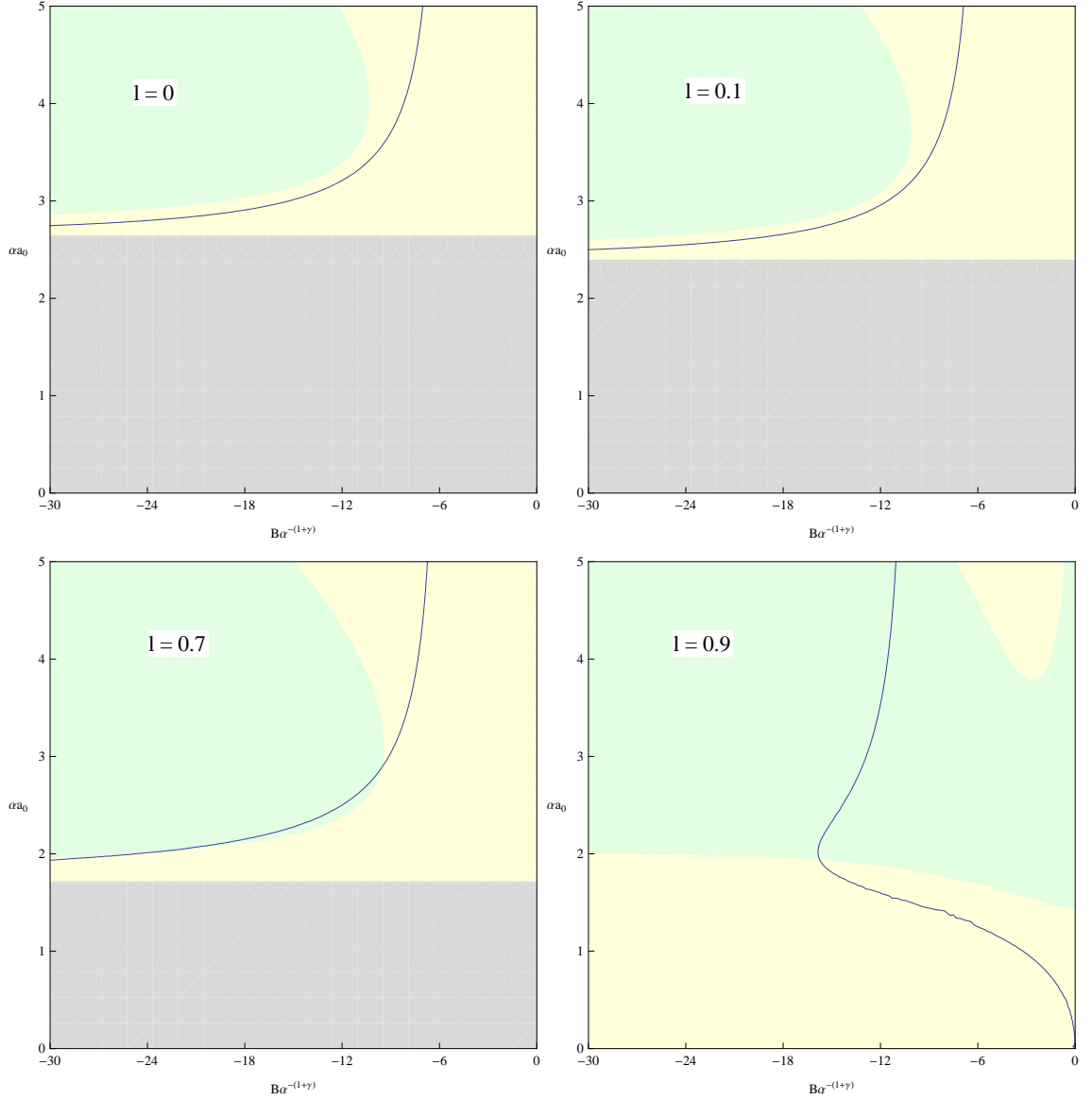


Figure 3: Plots for regular Hayward thin-shell wormholes by taking VDW quintessence with  $\gamma = -0.5$ ,  $M = 1$ ,  $\alpha = 1$  and different values of Hayward parameter  $l$ . The stable and unstable regions are represented by green and yellow colors, respectively, while the grey zone corresponds to non-physical region. Here  $B\alpha^{-(1+\gamma)}$  and  $\alpha a_0$  are labeled along abscissa and ordinate, respectively.

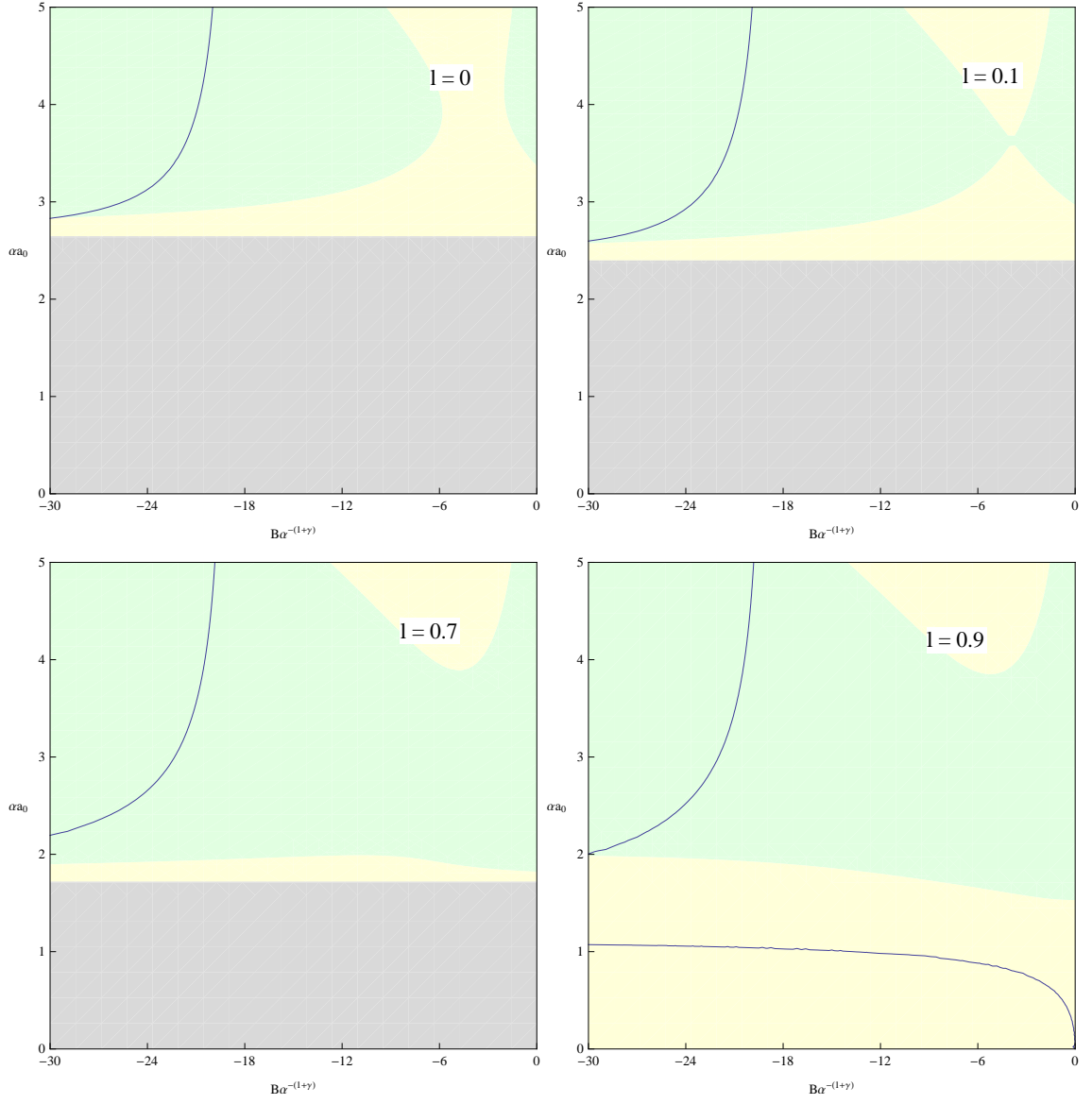


Figure 4: Plots for VDW quintessence EoS taking  $\gamma = 1$ ,  $M = 1$ ,  $\alpha = 1$  and different values of Hayward parameter  $l$ .

where  $E = \frac{B}{1+w} - 1$ ,  $B \in (-\infty, \infty)$ ,  $-C < w < 0$  and  $C$  is a positive constant rather than unity. The dynamical equation for static solutions through Eqs.(13) and (25) yield

$$\begin{aligned} & [a_0^2 F'(a_0) + 2a_0 F(a_0)][2a_0]^\gamma - 2(4\pi a_0^2)^{1+\gamma} [F(a_0)]^{\frac{1-\gamma}{2}} \\ & \times \left[ E + \left\{ (2\pi a_0)^{-(1+\gamma)} (F(a_0))^{\frac{(1+\gamma)}{2}} - E \right\}^{-w} \right] = 0. \end{aligned} \quad (26)$$

Differentiation of Eq.(25) with respect to  $a$  leads to

$$\sigma'(a) + 2p'(a) = \sigma'(a) \left[ 1 + 2w(1+\gamma) \{ \sigma^{1+\gamma} - E \}^{-(1+w)} - \frac{2\gamma p(a)}{\sigma(a)} \right], \quad (27)$$

which determines the second derivative of potential function as

$$\begin{aligned} \Psi''(a_0) &= F''(a_0) + \frac{(\gamma-1)[F'(a_0)]^2}{2F(a_0)} + \frac{F'(a_0)}{a_0} [1 + 2w(1+\gamma) \\ &\times \left\{ \left( \frac{\sqrt{F(a_0)}}{2\pi a_0} \right)^{1+\gamma} - E \right\}^{-(1+w)}] - \frac{2F(a_0)}{a_0^2} (1+\gamma) \\ &\times \left[ 1 + 2w \left\{ \left( \frac{\sqrt{F(a_0)}}{2\pi a_0} \right)^{1+\gamma} - E \right\}^{-(1+w)} \right]. \end{aligned} \quad (28)$$

Using Eq.(19) in (26), the corresponding dynamical equation for static solution becomes

$$\begin{aligned} & a_0^3 + 2Ml^2 - 4Ma_0^2 - 2(2\pi a_0)^{1+\gamma} [a_0^3 + 2Ml^2 - 2Ma_0^2]^{\frac{1-\gamma}{2}} [E \\ & + \{ (2\pi a_0)^{-(1+\gamma)} (a_0^3 + 2Ml^2)^{\frac{-(1+\gamma)}{2}} (a_0^3 + 2Ml^2 - 2Ma_0^2)^{\frac{(1+\gamma)}{2}} - E \}^{-w}] = 0, \end{aligned} \quad (29)$$

which gives static regular Hayward wormhole solutions. Here we again employ the same technique for the stability analysis as in the previous subsection. The results in Figures **5-6** correspond to GCCG. For  $\gamma = 0.2, 1$  and  $l = 0, 0.1, 0.7$ , we find fluctuating wormhole solutions. It is found that unstable solution exists for small values of throat radius  $a_0$ , while the solutions become stable when the throat radius expands. There exists a non-physical region for  $0 < a_0 < 2$  which diminishes with  $l = 0.9$  making a

stable traversable wormhole. For  $\gamma = 1$  and  $l = 0.9$ , there is a fluctuating behavior of wormhole throat. Initially, it is stable for smaller values of  $a_0$  but becomes unstable by increasing throat radius. Finally, we analyze stable regular Hayward thin-shell wormhole for throat radius  $a_0 > 2$  which undergoes expansion.

### 3.3 Modified Cosmic Chaplygin Gas

Here we assume MCCG model for exotic matter for which EoS is given by

$$p = A\sigma - \frac{1}{\sigma^\gamma} [E + (\sigma^{1+\gamma} - E)^{-w}]. \quad (30)$$

Sadeghi and Farahani [30] assumed MCCG as varying by considering  $E$  as a function of scale factor  $a$ , while in our case,  $E$  is assumed as a constant. Substituting Eq.(13) in (30), the dynamical equation (for static configuration) is

$$\begin{aligned} & [a_0^2 F'(a_0) + 2a_0 F(a_0)](1 + 2A)[2a_0]^\gamma - 2(4\pi a_0^2)^{1+\gamma} [F(a_0)]^{\frac{1-\gamma}{2}} \\ & \times \left[ E + \left\{ (2\pi a_0)^{-(1+\gamma)} (F(a_0))^{\frac{(1+\gamma)}{2}} - E \right\}^{-w} \right] = 0. \end{aligned} \quad (31)$$

The first derivative of EoS with respect to  $a$  yields

$$\sigma'(a) + 2p'(a) = \sigma'(a) \left[ 1 + 2(1 + \gamma) \{ A + w(\sigma^{1+\gamma} - E)^{-(1+w)} \} - \frac{2\gamma p(a)}{\sigma(a)} \right]. \quad (32)$$

It is noted that  $\Psi(a) = \Psi'(a) = 0$  at  $a = a_0$ , while the second derivative of  $\Psi(a)$  through Eq.(32) takes the form

$$\begin{aligned} \Psi''(a_0) &= F''(a_0) + \frac{(\gamma - 1)[F'(a_0)]^2}{2F(a_0)} + \frac{F'(a_0)}{a_0} [1 + 2 \{ A + w(1 + \gamma) \\ & \times \left\{ \left( \frac{\sqrt{F(a_0)}}{2\pi a_0} \right)^{1+\gamma} - E \right\}^{-(1+w)} \}] - \frac{2F(a_0)}{a_0^2} (1 + \gamma) \\ & \times \left[ 1 + 2(1 + \gamma) \left[ A + w \left\{ \left( \frac{\sqrt{F(a_0)}}{2\pi a_0} \right)^{1+\gamma} - E \right\}^{-(1+w)} \right] \right] \end{aligned} \quad (33)$$

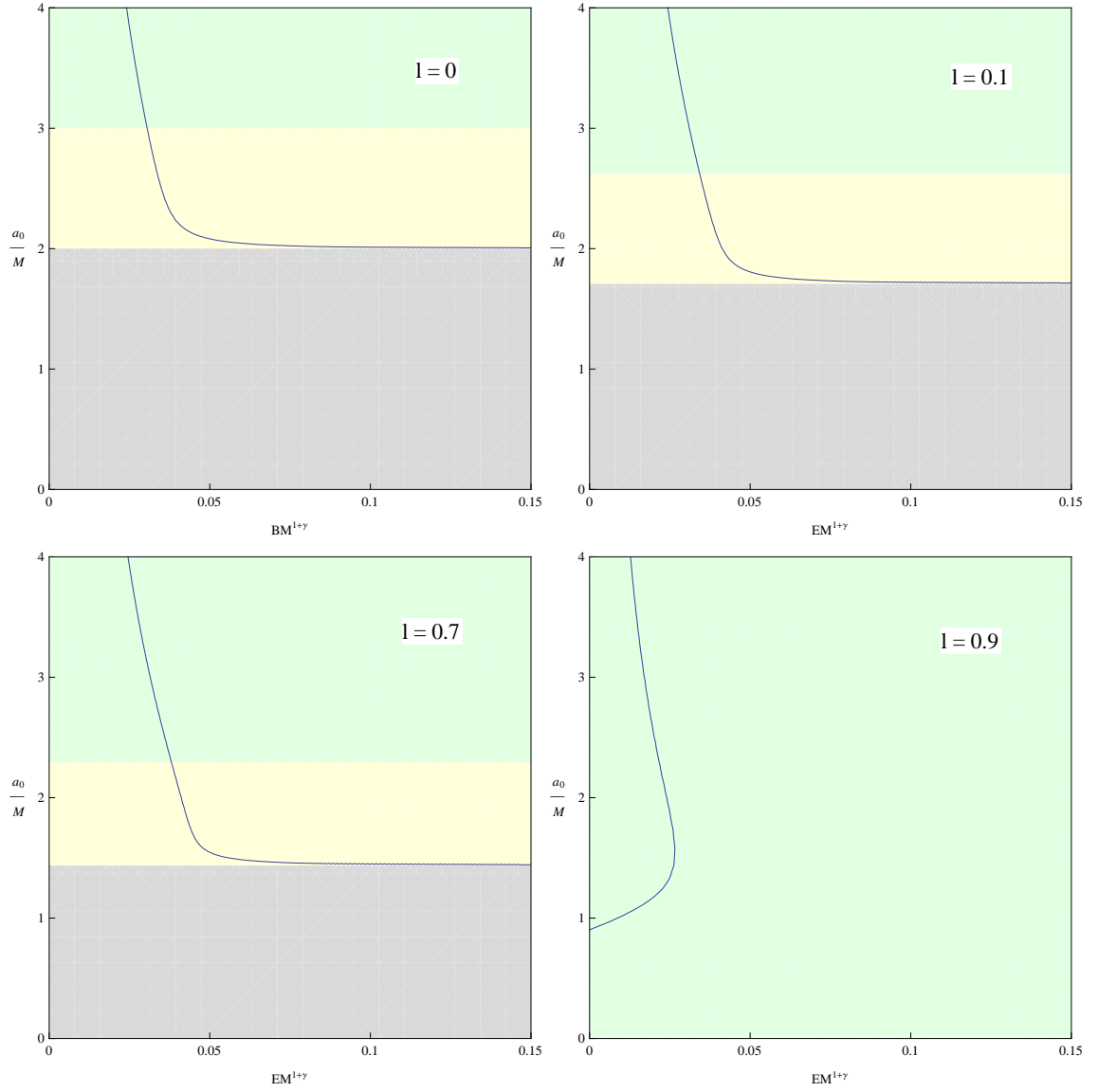


Figure 5: Plots for GCCG with parameters  $\gamma = 0.2$ ,  $M = 1$ ,  $w = -10$  and different values of  $l$ . Here  $EM^{(1+\gamma)}$  and  $\frac{a_0}{M}$  are labeled along abscissa and ordinate, respectively.



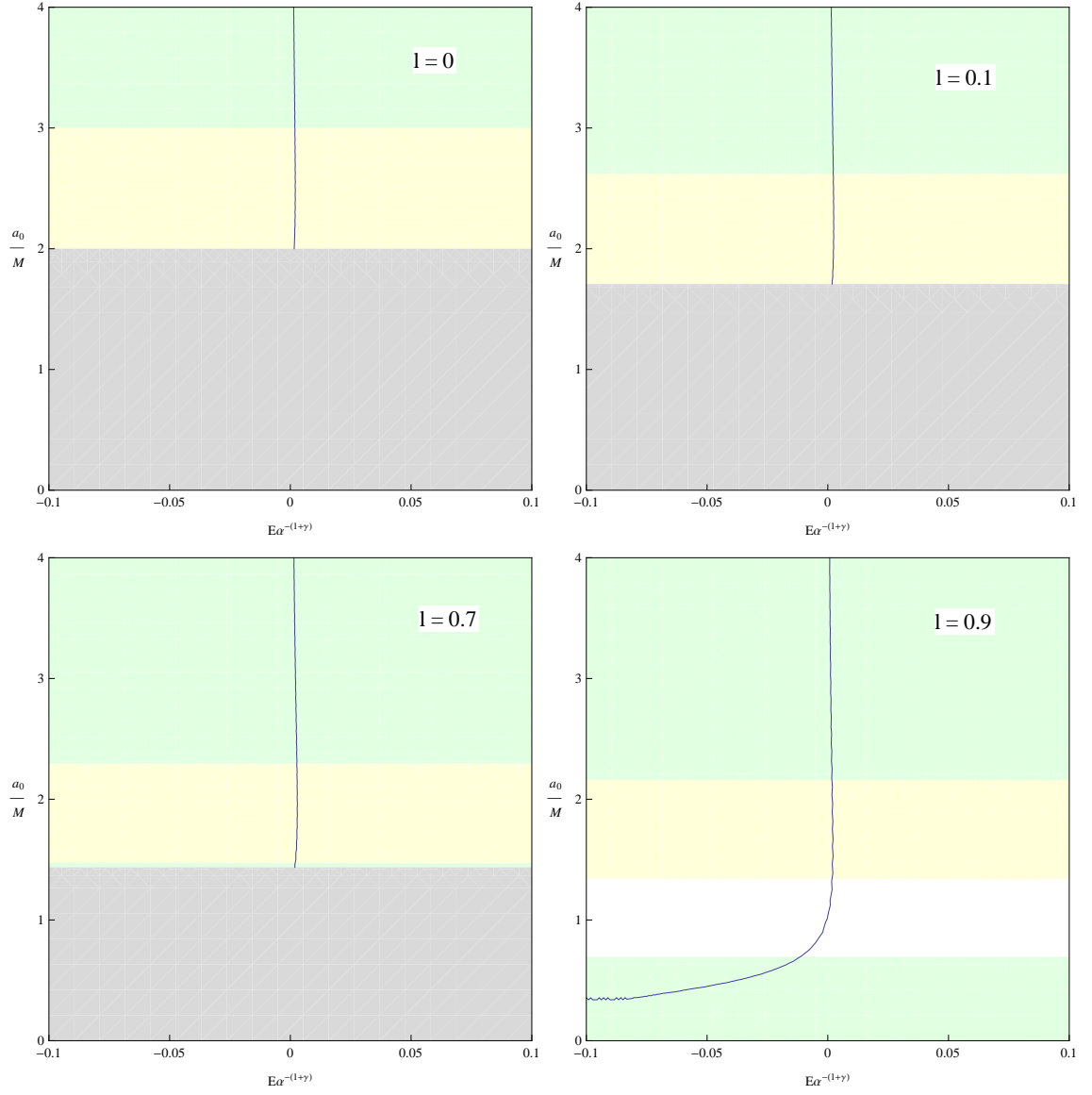


Figure 6: Plots for GCCG with  $\gamma = 1$ ,  $M = 1$ ,  $w = -10$  and different values of  $l$ .

For Hayward wormhole static solutions, Eq.(31) turns out to be

$$\begin{aligned}
& a_0^3(a_0^3 + Ml^2 - 4Ma_0^2) + 4M^2l^2(l^2 - 4a_0^2) + 2A - (2\pi a_0)^{1+\gamma} \\
& \times (a_0^3 + 2Ml^2)^{\frac{3+\gamma}{2}}(a_0^3 + 2Ml^2 - 2Ma_0^2)^{\frac{-(1+\gamma)}{2}} \left[ E + \{(2\pi a_0)^{-(1+\gamma)} \right. \\
& \times (a_0^3 + 2Ml^2)^{\frac{-(1+\gamma)}{2}}(a_0^3 + 2Ml^2 - 2Ma_0^2)^{\frac{(1+\gamma)}{2}} - E \}^{-w} \Big] = 0. \quad (34)
\end{aligned}$$

The results in Figures 7-8 show that for  $l = 0, 0.1, 0.7$ , one stable as well as unstable solution exist for  $\gamma = 0.2, 0.6$ . We analyze that two stable and two unstable regions appear for  $l = 0.9$ . There exists a non-physical region for  $0 < a_0 < 2$  again showing the event horizon which continues to decrease and eventually vanishes for  $l = 0.9$  making a traversable wormhole. For  $l = 0.9$ , we find a traversable wormhole solution with fluctuating behavior of wormhole throat which shows a stable wormhole solution with throat expansion. It is noted that the stability region increases by increasing values of  $l$ .

## 4 Concluding Remarks

In this paper, we have constructed regular Hayward thin-shell wormholes by implementing the Visser's cut and paste technique and analyzed their stability by incorporating the effects of increasing values of Hayward parameter. The surface stresses have been found by using Lanczos equations. The sum of surface stresses of matter indicates the violation of NEC which is a fundamental ingredient in wormhole physics leading to the existence of exotic matter. It is found that the wormhole has attractive or repulsive characteristics corresponding to  $a^r > 0$  and  $a^r < 0$ , respectively. For a convenient trip through wormhole, an observer should not be dragged away by enormous tidal forces which requires that the accelerations felt by observer must not exceed the Earth's acceleration. The construction of viable thin-shell wormholes depends on total amount of exotic matter confined within the shell. Here we have quantified the total amount of exotic matter by the volume integral theorem which is consistent with the fact that a small quantity of exotic matter is needed to support the wormhole. We have obtained a dynamical equation which determines possible throat radii for the static wormhole configurations.

It is found that one may construct a traversable wormhole theoretically with arbitrarily small amount of fluids describing cosmic expansion. In this

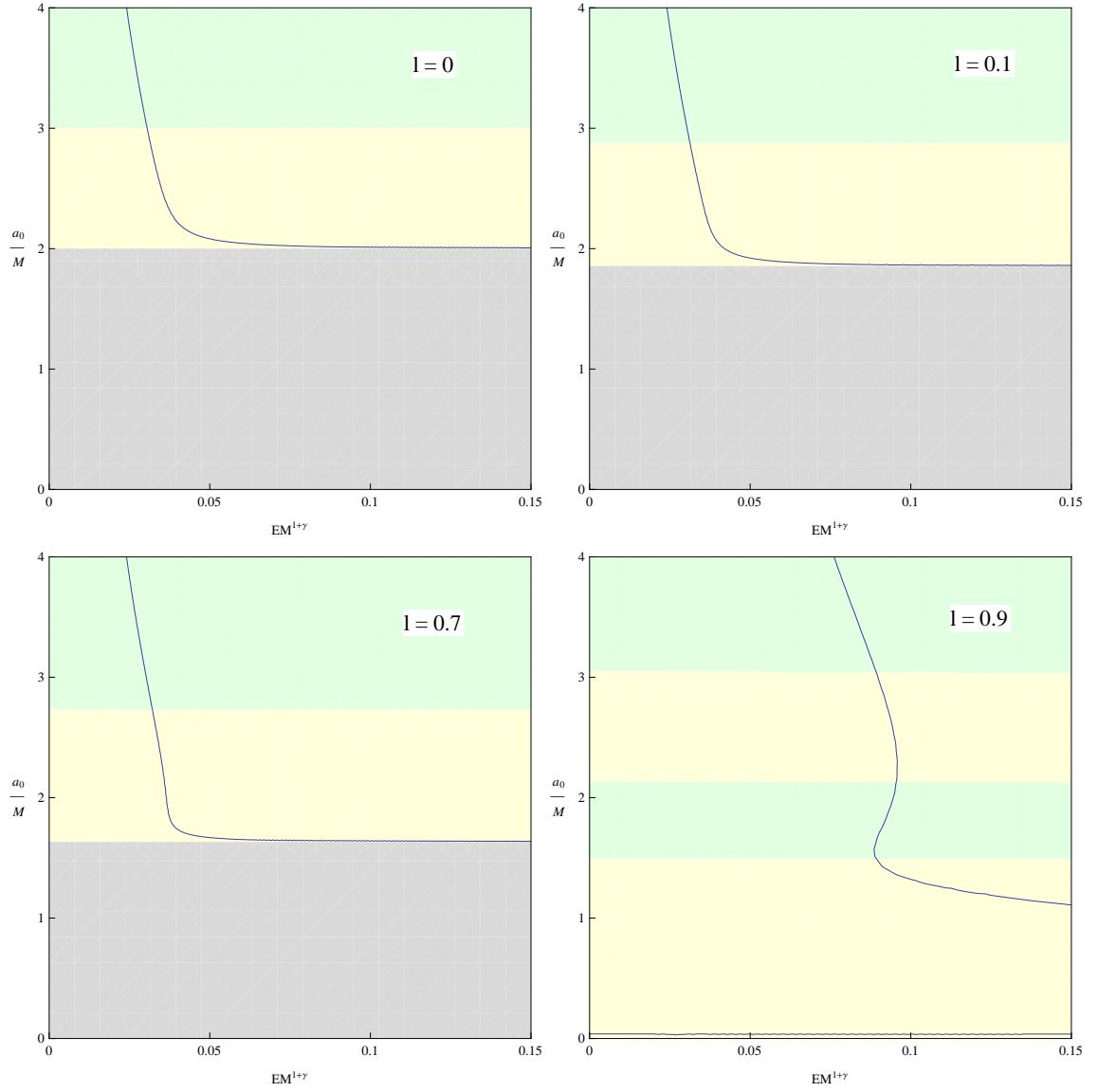


Figure 7: Plots for MCCG with  $\gamma = 0.2$ ,  $M = 1$ ,  $A = 1$  and  $w = -10$  with different values of  $l$ . Here  $EM^{(1+\gamma)}$  and  $\frac{a_0}{M}$  are labeled along abscissa and ordinate, respectively.

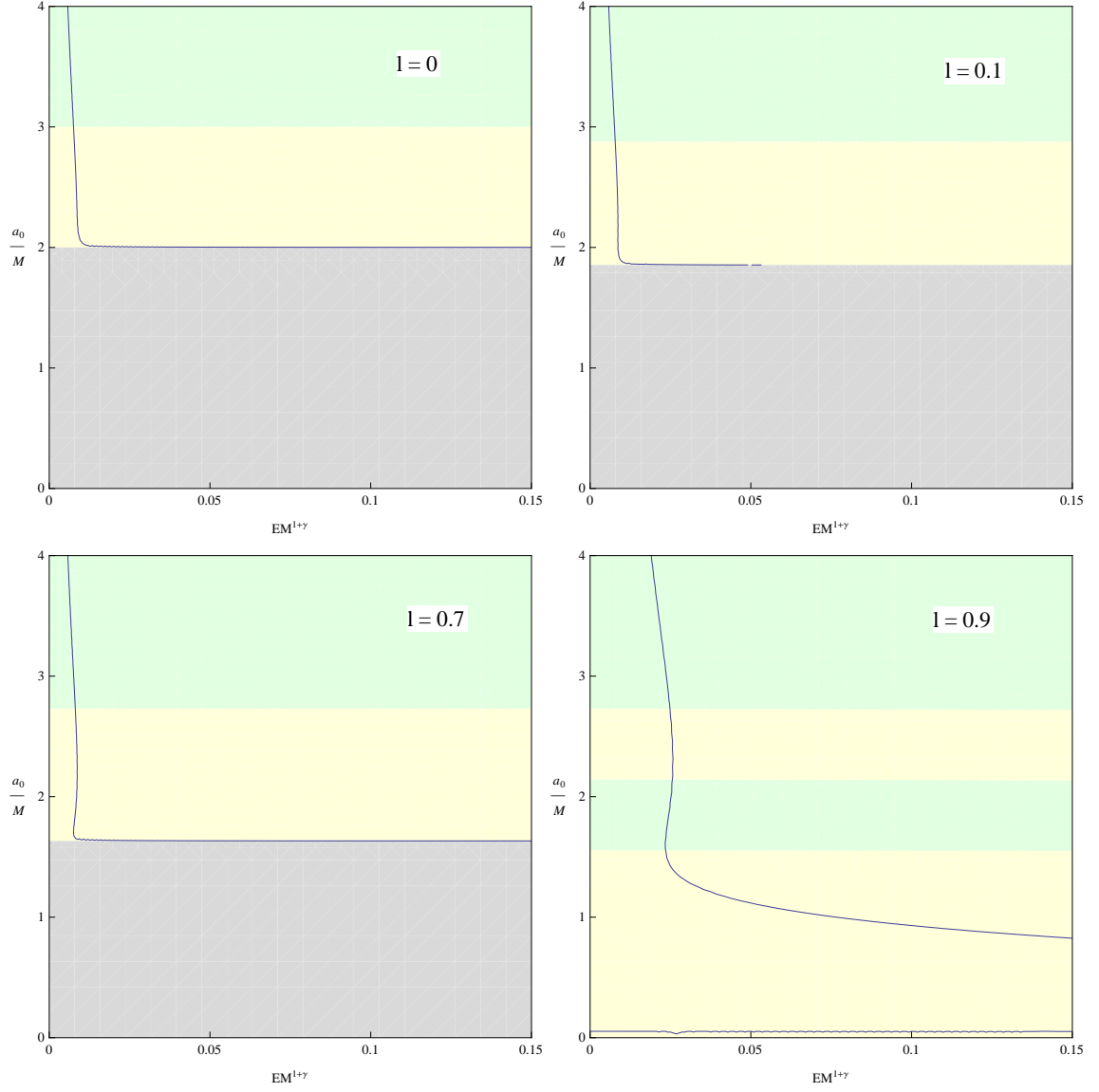


Figure 8: Plots corresponding to MCGG with parameters  $\gamma = 0.6$ ,  $M = 1$ ,  $A = 1$ ,  $w = -10$  and different values of  $l$ .

context, we have analyzed stability of the regular Hayward thin-shell wormholes by taking VDW quintessence, GCCG and MCCG models at the wormhole throat. Firstly, we have investigated the possibility of stable traversable wormhole solutions using VDW quintessence EoS which describes cosmic expansion without the presence of exotic fluids. We have analyzed both stable and unstable wormhole configurations for small values of EoS parameter  $\gamma$ . The graphical analysis shows a non-physical region (grey zone) for  $0 < a_0 < 2.8$  (Figures 3-4). In this region, the stress-energy tensor may vanish leading to an event horizon and the wormhole becomes non-traversable producing a black hole. The non-physical region in the wormhole configuration decreases gradually and vanishes for  $l = 0.9$ . In this case, we find unstable wormhole solution as  $B\alpha^{-(1+\gamma)}$  approaches to its maximum value. We have also examined stability of Hayward thin-shell wormholes for  $\gamma \in [1, \infty)$  and found only stable solutions with  $l = 0, 0.1, 0.7$  which show expansion of the wormhole throat. It is worth mentioning here that regular Hayward thin-shell wormholes can be made traversable as well as stable by tuning the Hayward parameter to its large value. Also, it is noted that VDW quintessence fluid minimizes the usage of exotic matter and  $\gamma = 1$  is the best fitted value which induces only stable solutions.

In case of GCCG, we have analyzed fluctuating (stable and unstable) solutions for  $l = 0, 0.1, 0.7$  and  $\gamma = 0.2, 1$ . It is found that unstable solution exists for small values of throat radius  $a_0$ , which becomes stable when the throat radius expands. There exists a non-physical region for  $0 < a_0 < 2$  which diminishes for  $l = 0.9$  making a stable traversable wormhole. For  $\gamma = 1$ , we have analyzed stable regular Hayward thin-shell wormhole for throat radius  $a_0 > 2$  which undergoes expansion. We have found that only stable wormhole configurations exist for the Hayward parameter  $l = 0.9$ . Finally, for MCCG, we have found one stable and one unstable region for  $l = 0, 0.1, 0.7$ , while these regions become double (two stable and two unstable) for  $l = 0.9$  (Figures 7-8). There exists a non-physical region for  $0 < a_0 < 2$  again showing the event horizon which eventually vanishes for  $l = 0.9$  making a traversable wormhole. It is noted that the stability region increases by increasing values of  $l$ . It was found that small radial perturbations yield no stable solutions for the regular Hayward thin-shell wormholes [27]. We conclude that stable regular Hayward wormhole solutions are possible against radial perturbations for VDW quintessence, GCCG and MCCG models. The trivial case  $l = 0$  corresponds to the Schwarzschild thin-shell wormhole. It is worth mentioning here that the Hayward parameter increases stable regions for regular Hay-

ward thin-shell wormholes.

**The authors have no conflict of interest.**

## References

- [1] Misner, C.W. and Wheeler, J.A.: Ann. Phys. **2**(1957)525.
- [2] Hayward, S.A.: arXiv gr-qc/0203051.
- [3] Morris, M. and Thorne, K.: Am. J. Phys. **56**(1988)395.
- [4] Lemos, J.P.S.: Phys. Lett. B **352**(1995)46.
- [5] Israel, W.: Nuovo Cimento B **B48**(1967)463.
- [6] Nandi, K.K., Zhang, Y.Z. and Kumar, K.B.V.: Phys. Rev. D **70**(2004)127503.
- [7] Visser, M.: Phys. Rev. D **39**(1989)3182; *Lorentzian Wormholes*, (AIP Press, 1996).
- [8] Kim, S.W. and Lee, H.: Phys. Rev. D **63**(2001)064014.
- [9] Ishak, M. and Lake, K.: Phys. Rev. D **65**(2002)044011.
- [10] Lobo, F.S.N. and Crawford, P.: Class. Quantum Grav. **21**(2004)391.
- [11] Eiroa, E.F. and Romero, G.E.: Gen. Relativ. Gravit. **36**(2004)4.
- [12] Sharif, M. and Azam, M.: J. Cosmol. Astropart. Phys. **05**(2013)025.
- [13] Sharif, M. and Muntaz, S.: arXiv 1604.01012.
- [14] Sharif, M. and Muntaz, S.: Astrophys. Space Sci. **361**(2016)218.
- [15] Das, A. and Kar, S.: Class. Quantum Grav. **22**(2005)3045.
- [16] Lobo, F.S.N.: Phys. Rev. D **73**(2006)064028.
- [17] Eiroa, E.F. and Simeone, C.: Phys. Rev. D **76**(2007)024021.
- [18] Lobo, F.S.N.: Phys. Rev. D **80**(2009)044033.

- [19] Capozziello *et al.*: J. Cosmol. Astropart. Phys. **04**(2005)005.
- [20] Eiroa, E.F.: Phys. Rev. D **80**(2009)044033.
- [21] Kuhfittig, P.K.F.: Acta Phys. Polo. B **41**(2010)09.
- [22] Sharif, M. and Azam, M.: Eur. Phys. J. C **73**(2013)2554.
- [23] Sharif, M. and Mumtaz, S.: Adv. High Energy Phys. **2014**(2014)13.
- [24] Lobo, F.S.N.: Phys. Rev. D **75**(2007)024023.
- [25] Sharif, M. and Mumtaz, S.: Astrophys. Space Sci. **352**(2014)729.
- [26] Sharif, M. and Mumtaz, S.: Can. J. Phys. **93**(2015)12.
- [27] Halilsoy, M., Ovgun, A. and Mazharimousavi, S.H.: Eur. Phys. J. C **74**(2014)2796.
- [28] Hayward, S.A.: Phys. Rev. Lett. **96**(2006)031103.
- [29] Gonzalez-Diaz, P.F.: Phys. Rev. D **68**(2003)021303.
- [30] Sadeghi, J. and Farahani, H.: Astrophys. Space Sci. **347**(2013)209.